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ENERGETICS AND INTERPLANETARY EFFECTS OF THE AUGUST 14
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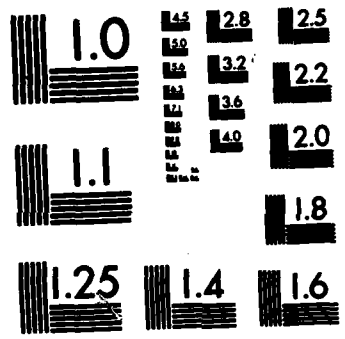
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AU as $\sim R^{-0.8}$

approximately
(3.14)
greater than or equal to
to the minus 0.8 power

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SPACE SCIENCES LABORATORY

ENERGETICS AND INTERPLANETARY EFFECTS

OF THE AUGUST 14 AND 18, 1979 SOLAR FLARES

A

Summary of observations made during the SMY/STIP Event No. 1:
August 14-18, 1979

Authors: S. R. Kane, M. K. Bird, V. Domingo, G. Green,
G. R. Gapper, A. Hewish, R. A. Howard, B. Iwers, B. V. Jackson,
U. Koren, H. Kunow, R. Muller-Mellin, B. Rompolt, B. Sanahuja,
H. S. Sawant, R. T. Stewart, G. Wibberenz, and P. Zlobec

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Energetics and Interplanetary Effects of the August 14 and 18, 1979 Solar Flares

Summary of observations made during the SMY/STIP Event No. 1:

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ABSTRACT

The STIP Event No. 1, which covered the time interval August 14 - 18, 1979, was characterized by two energetic flares: one on August 14 (~ 1243 UT) and the other on August 18 (~ 1400 UT). The hard X-ray, soft X-ray, optical, radio and energetic particle emissions from these flares and their interplanetary effects were observed with many instruments in space and on the ground. A summary of some of these observations is presented. The results of a preliminary analysis relevant to the acceleration of particles, coronal transients and evolution of shocks are as follows: (1) During the August 14 flare energetic particles were probably accelerated but could not escape the Sun in large numbers. On the other hand, during the August 18 flare the acceleration of high energy particles occurred relatively high in the corona, from whence they could easily escape into interplanetary space but could not penetrate down to the lower altitudes in the solar atmosphere in large numbers. (2) The kinetic energy of the coronal transient associated with the August 14 flare was much larger than the total energy of energetic electrons, indicating an additional energy source for the transient. (3) The shock associated with the August 18 flare extended to $\geq 2 \pi$ steradians. The shock maintained its speed from the



flare site to a distance of $\sim 35 R_{\odot}$ and then decelerated to a distance of ~ 1 AU as $\sim R^{-0.5}$.

I. Introduction

Energetic solar flares are known to produce a variety of electromagnetic emissions extending from gamma-ray energies to radio wavelengths. The flares also produce mass motions, shock waves and energetic particles detectable in interplanetary space. The terrestrial effects of such flares vary from short wave fadeout and polar cap absorption to increases in the neutron monitors located at high latitudes. The primary energy source for all these phenomena is believed to be the magnetic field (or equivalent electric currents) in the vicinity of the optical flare. The release of energy seems to occur in the form of the acceleration of particles and/or the heating of the ambient plasma.

Although the above general aspects of the flares have now been well observed, the relevant physical processes at the Sun, as well as in interplanetary space, are not well understood. For example, it is not clear whether the acceleration of particles occurs in one or more stages or what role the shocks play in the acceleration process. The distribution of energy dissipated in different processes such as energetic particles, heating of plasma, mass motions, shocks, etc. is also not known. The generation of shocks and their evolution as they propagate from the Sun is another aspect of the flare phenomenon which is very little understood.

The study of solar flares and their interplanetary effects is often made difficult by the absence of a well-coordinated, wide range of observations for individual flares. This difficulty was partially overcome during the SMY/STIP Event No. 1 when institutions and individuals in many different countries participated in a joint observational effort. This report presents a summary of those observations and the progress made so far in their analysis and interpretation. It is hoped that a detailed analysis of this STIP event will add considerably to our understanding of the acceleration and propagation of energetic particles, the formation and evolution of shocks, and other phenomena associated with large (energetic) solar flares.

II. General Characteristics of STIP Event No. 1

The period August 14 - 18, 1979, covered by STIP Event No. 1, was characterized by two energetic flares: one on August 14 (~ 1243 UT) and the other on August 18 (~ 1400 UT). Both flares produced large hard X-ray bursts and other energetic phenomena, such as energetic particles, coronal transients, and shock waves, although some phenomena may not have been observed because of the locations of the flares near the east limb of the Sun and/or data gaps in the observations. In the August 18, 1979, flare, the peak ~ 30 keV X-ray flux was as large as that from the well-known large flare on August 4, 1972. Although the peak ~ 30 keV X-ray flux in the August 14, 1979 flare was smaller than that in the August 4, 1972 flare by a factor of ~ 10, the hardness of the X-ray spectrum and characteristic features such as the magnitude and variation of the gradual hard X-ray emission were similar in the two flares. Thus STIP Event No. 1 consists of two energetic flares comparable in some respects to the very large August 4, 1972 flare.

In addition to the ground-based instruments and near-Earth spacecraft, a number of interplanetary spacecraft were also in operation during this period (cf. Figure 1; *Vette*, 1981). They provide a variety of observational information which is complementary/supplementary to that obtained with the near-Earth instruments. We present below the observations of the two August 1979 flares in some detail and then discuss briefly some of the implications of these observations.

III. 14 August 1979 Flare

The principal optical and radio characteristics of the 14 August 1979 flare are summarized in Table 1. The optical flare was of importance 1B and was located at S22,E73 in McMath Region 16224. The flare started at ~ 1240 UT and reached two successive maxima at ~ 1244 UT and ~ 1251 UT followed by a long decay during the next two or more hours.

Prior to the optical flare, a large and unusually bright eruptive prominence was observed in H α by the Wroclaw Astronomical Observatory (Poland) at S32-E from 1140 to 1216 UT

[Rompolt, 1982]. The prominence developed into an expanding arch in the later phase of evolution. The H α material was observed up to a height of 3×10^5 km above the limb. The velocity of eruption for the top part of the expanding arch varied from 120 km/sec to 285 km/sec during the observational period. The synoptic maps indicate that a twisted filament was present earlier in the place of the erupting prominence.

Figure 2 [Vilmer *et al.*, 1982] shows the flare-associated hard X-ray burst observed with the ISEE-3 spacecraft. The spectrum averaged over the whole event is shown in Figure 3. The impulsive and gradual hard X-ray emissions can be very clearly identified in this event. Spectra of both types of emissions are relatively hard. During the gradual phase all X-ray energies do not peak at the same time but exhibit an energy dependent delay, the higher energy X-rays reaching their maxima later. This behavior has been observed before in only three other large flares, viz., those on 4 August 1972, 7 August 1972, and 22 November 1977.

The microwave emission recorded at Segamore Hill Observatory is shown in Figure 4 [Sawant and Kane, 1982]. The impulsive maximum was not well developed, as can be seen from the Ottawa measurements (Figure 5). The two gradual maxima are shown separately in Figure 4. The gradual emission extended to ~ 35 GHz frequency and exhibited a frequency-dependent delay in the time of maximum, the highest frequency emission reaching the maximum earliest.

Intense decimetric and metric type IV bursts starting at ~ 1243 UT have been recorded by the Dwingeloo and Weissenau radiospectrographs, respectively (Solar Geophysical Data). The observations of 237 - 505 MHz type IV emission made at the Trieste Astronomical Observatory [Zlobec and Koren, 1982] show an unusually dense and strong spike activity superimposed on the smoothly varying time profile. The continuum was unpolarized. However, unlike most type IV events, the polarization (right-handed) of spikes was evolving in time. No type II bursts in close association with the flare have been reported. A weak SA (shock-accelerated) event [Cane *et al.*, 1981] was observed by the ISEE-3, however, indicating the presence of a weak shock and acceleration of electrons [Cane, 1982].

There is no clear evidence for the release of a substantial number of energetic particles from this flare. Measurements of energetic particles are available from Helios-1 and -2 (Figure 6 and Figure 7; *Kunow et al.*, 1982) and IMP-8 and ISEE-3 spacecraft [*McGuire et al.*, 1982; *Lin*, 1982]. Low energy particle measurements are available from ISEE-3 (Figure 8; *Domingo and Sanahuja*, 1982). A few hours after the 14 August flare there was a small increase in the low energy proton flux. At that time (~ 2100 UT), the flux associated with the previous event was still quite large. The timing of these particles is compatible with near simultaneous injection of particles of all energies in the corona. Therefore it is plausible that they have been accelerated near the Sun by the shock produced by the flare.

On August 17, between 0925 and 1015 UT (data gap), a new particle flux increase started at all energies. The maximum was reached on August 17. The association of these particles with the August 14 flare is uncertain because of the occurrence of another 1B flare on August 16 at N18,E25.

The "Solwind" white light coronagraph aboard the P78-1 satellite recorded a coronal transient at ~ 1337 UT on August 14 1979 (Figure 9; *Poland et al.*, 1981). The transient was located on the SE limb (20°) and was one of the most massive ($\sim 2 \times 10^{16}$ g) and fastest (~ 900 km sec $^{-1}$) transients ever observed.

Interplanetary scintillation (IPS) observations have been made at the University of California, San Diego and in Cambridge, England following the 14 August flare (Figure 10, *Jackson*, 1982a; Figure 11, *Jackson, Hewish, and Gapper*, 1982). Sky maps of IPS made at Cambridge show a major enhancement during Aug. 15-22 of the kind usually associated with a corotating stream. The disturbance approached from the East and passed over the Earth on Aug. 19-20. This was probably caused by the fast solar wind stream which was observed on ISEE-3 during Aug. 20-21. No clear evidence was found for expanding spherical shells which might have been associated with either of the flares on Aug. 14 and 18. [*Hewish and Gapper*, 1982].

A composite of NRL, San Diego, and Cambridge data from *Jackson* [1982b] is given in Figure 12. Lines dashed and solid give elongations of two simple constant velocity models of

the position of the outermost extent of the mass of the ejection. Neither model fits the data well. *Jackson* [1982a] proposes that the material of this mass ejection extends over a large range of longitude ($>90^\circ$) as well as latitude throughout the interplanetary medium. The radial speed of the ejection is not uniform with solar longitude, however, and especially, is considerably slower ($\sim 400 \text{ km sec}^{-1}$) in the direction towards earth along the ecliptic plane.

Other sources of information regarding the solar wind plasma are the IMP-8 and Pioneer-12 instruments (Figure 13, and Figure 14; Solar Geophysical Data) and ISEE-3 Plasma Waves instrument [*Scarf*, 1981]. There appears to be no detection of the shock near the Earth or a geomagnetic sudden commencement following the 14 August 1979 flare.

IV. 18 August 1979 Flare

Figure 15 shows the hard X-ray emission observed with ISEE-3 during the 18 August 1979 flare [*Kane*, 1981]. This figure and Table 2 also present a summary of other emissions, such as $H\alpha$, soft X-rays (SMS-GOES) and radio, obtained from Solar Geophysical Data.

The optical flare (~ 1420 UT, -B, N08,E90) often associated with this event occurred later than the largest maximum in the hard X-ray burst. This flare is therefore not directly related to the main event. It appears that two relatively large flares occurred in the same active region considerably before this (~ 1420 UT) flare. Those two flares probably occurred somewhat behind the East limb and so could not be observed at optical wavelengths. The hard X-ray, soft X-ray, and microwave radio measurements are consistent with this inference.

The ~ 30 keV X-ray flux at the large maximum (~ 1420 UT) is the largest X-ray flux observed by ISEE-3 in the 1978-79 period and is comparable to the very large flare on 4 August 1972. The X-ray spectrum (Figure 16), however, is considerably softer in the 18 August 1979 flare. It is possible that part of the hard X-ray source was occulted from the line of sight of the ISEE-3 instrument.

The microwave emission from this flare, which was recorded at Ottawa, Boulder, and other observatories, had three pronounced maxima as in the hard X-ray emission (Figure 5;

Solar Geophysical Data). However, the peak flux (490 SFU at 2.8 GHz) is smaller than that in the 14 August 1979 flare (4030 SFU at 2.8 GHz) by a factor of ~ 8 . Although this is consistent with the softer hard X-ray spectrum, it is surprising that an apparently energetic flare produced a very modest microwave flux. It is possible that part of the microwave source was occulted from the line of sight.

Decimetric (type III-G-Int. 2) and metric (type IV-Int. 3) radio bursts have been reported by the Dwingeloo and Weissenau radiospectrographs, respectively. Measurements of the type IV emission made at the Trieste Astronomical Observatory [Zlobec and Koren, 1982] indicate rapid pulsations in the 237 MHz emission (left-handed channel). The intensity of the continuum decreased with increasing frequency and the polarization was 0%. A metric type II burst was recorded at Fort Davis from 1412 to 1433 UT [Maxwell, 1982]. A type II burst was also observed by ISEE-3 in the KHz frequency range (Figure 17; Cane *et al.*, 1982).

Energetic particles (electrons, protons and He nuclei) were observed by Helios-1 and -2 (Figures 6 and 7 Kunow *et al.*, 1982) and IMP-8 instruments (Figure 18; McGuire *et al.*, 1982). Protons with energy ≥ 50 MeV were detected on August 19. Faster rise time and larger peak flux recorded by Helios-2 compared with Helios-1 indicates that the former spacecraft was magnetically better connected with the flare site, and hence measured the "true magnitude" of the particle event. This is consistent with the locations of the two spacecraft relative to the Sun (Figure 19; based on Kunow *et al.*, 1982). The proton-to-alpha ratios are, however, comparable at the two spacecraft. The modulation of the particle flux during the decay is probably caused by interplanetary shocks. This is difficult to verify, however, since no plasma or magnetic data are presently available from either Helios spacecraft.

Low energy particle measurements from ISEE-3 are also available (Figure 8; Domingo, 1981). The relatively slowly rising flux is consistent with the location of the flare near the east limb of the Sun.

The flare produced a shock which was well observed near the east limb (1500 UT, 18 August) at a distance of $\sim 13 R_{\odot}$ from the Sun (Figure 20; Woo and Armstrong, 1981) and also

near the Earth (0550 UT, 20 August) at the ISEE-3 location (Figure 21; *Smith, 1982; Scarf, 1981*). A geomagnetic sudden commencement was recorded at ~ 0600 UT on August 20. Enhancement in the solar wind parameters was also recorded by IMP-8 (Figure 13; Solar Geophysical Data). Thus the shock seems to have extended to $\geq 2\pi$ steradians. Using the ground-based scattering observations of the 2.3 and 8.4 GHz radio signals from the Voyager 1 spacecraft, Woo and Armstrong have derived the solar wind flow speed and characteristic parameters of the turbulence such as electron density fluctuations. They found a shock speed of 3500 km sec^{-1} at $13.1 R_{\odot}$ and an average shock speed of 3509 km sec^{-1} from the Sun to $13.1 R_{\odot}$. A more appropriate value for this average shock speed has been estimated to be $\sim 2350 \text{ km sec}^{-1}$ [*Maxwell and Dyer, 1982; Woo, 1982*]. The ISEE-3 measurements [*Smith, 1982*] indicate an average shock speed of 1053 km sec^{-1} between the Sun and $\sim 0.99 \text{ AU}$ and a speed of 875 km sec^{-1} between $\sim 0.99 \text{ AU}$ and $\sim 1 \text{ AU}$. *Cane et al., [1982]* have examined the kilometric type II emission observed with the ISEE-3 spacecraft and other available data (Figure 22 and Figure 23). They have found that the shock speed was constant or possibly increased up to a distance of $35 R_{\odot}$ and then decreased at larger distances from the Sun as $\sim R^{-0.3}$.

Observations of coronal transients are not available for this event.

V. Summary and Discussion

The STIP Event No. 1 has provided us two relatively large flares which are remarkable in many ways. We discuss below some of these distinctive features:

A. Acceleration of Particles at the Sun

The August 14 flare had all the characteristics of an energetic flare. Although the optical flare and soft X-ray burst were of medium magnitude, the hard X-ray, microwave, and metric type IV radio bursts were intense. Both impulsive and gradual components were present in the hard X-ray emission. The total energy of electrons $\geq 25 \text{ keV}$ deduced from the hard X-ray emission produced in a thick target chromospheric source was $\sim 3 \times 10^{30}$ ergs. The long dura-

tion and hard spectrum of the gradual hard X-ray component extending into MeV range is often considered as an indication of the second stage of particle acceleration. However, no metric type II burst has been reported. Also, no significant fluxes of energetic particles were detected in interplanetary space. This could be due to any or all of the following: (a) unfavorable location of the flare — E73, (b) relatively closed magnetic field structure near the site of the particle acceleration region, or (c) relatively low efficiency of the particle acceleration process during this flare. It seems very likely that particles were accelerated during the August 14 flare but could not escape the Sun in sufficiently large numbers.

On the other hand, the August 18 flare produced energetic electrons, protons and heavier nuclei, which were detected in interplanetary space. The type II radio burst and interplanetary shock were well observed. The metric type IV radio burst (intensity 3) and the soft X-ray burst (X 6) were also very intense. The hard X-ray emission had a large flux at ~ 30 keV, the total electron energy ≥ 25 keV being $\sim 4 \times 10^{32}$ ergs. However, the spectrum for the impulsive as well as the gradual hard X-ray components was rather steep. There was no detectable flux of X-rays ≥ 400 keV. The microwave emission was also relatively weak. Thus, in the hard X-ray and microwave emissions, there is no clear signature of the second stage of particle acceleration. It appears that, in this flare, the acceleration of high energy particles occurred relatively high in the corona, from whence they could escape easily into interplanetary space but could not penetrate down to the lower altitudes in the solar atmosphere in sufficiently large numbers.

(b) Coronal Transients

The massive coronal transient observed in association with the August 14 flare corresponds to a kinetic energy of $\geq 4 \times 10^{31}$ ergs if $\geq \frac{1}{2}$ the mass of $\sim 2 \times 10^{16}$ g was moving at a speed of ~ 900 km sec $^{-1}$. This energy is at least ~ 10 times larger than the total energy ($\sim 3 \times 10^{30}$ ergs) in ≥ 25 keV electrons deduced from hard X-ray emission produced in a chromospheric thick target bremsstrahlung source. This energy can be increased by a factor of ~ 2 if electrons down to ~ 10 keV are included. Also, if the X-ray source for the gradual hard X-ray emission is coronal rather than chromospheric, the total energy carried by the energetic elec-

trons will be considerably higher. However, there is an indication here that the kinetic energy of the coronal transient was larger than that carried by energetic electrons, indicating an additional source of energy in the flare. This is consistent with the observations of coronal transients even in the absence of large solar flares [Poland *et al.*, 1981].

(c) Interplanetary Shock

The interplanetary shock associated with the August 18 flare seems to have extended over $\geq 2\pi$ steradians. It appears to have maintained its speed or accelerated from the flare site up to a distance of $\sim 35 R_{\odot}$ and then decelerated out to a distance of ~ 1 AU as $\sim R^{-0.8}$. It has been suggested by Cane *et al.* [1982] that the shock was driven out to $\sim 35 R_{\odot}$ and then was a blast wave out to a distance of ≥ 1 AU.

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Figure Captions

- Figure 1. Locations of interplanetary spacecraft during STIP Event No. 1 [Vene, 1981].
- Figure 2. Hard X-ray emission observed by ISEE-3 in association with the August 14, 1979 flare [Vilmer *et al.*, 1982]. Information about the optical flare, soft X-rays and radio bursts is also given.
- Figure 3. ISEE-3 observations of the X-ray spectrum averaged over the entire hard X-ray burst associated with the 14 August 1979 flare [Kane, 1982].
- Figure 4. 2.7 - 3.5 GHz microwave emission observed at Sagamore Hill at the time of the August 14, 1979 flare [Sawant and Kane, 1982].
- Figure 5. 2.8 GHz emission observed at Ottawa during the August 14 and 18, 1979 flare [Solar Geophysical Data].
- Figure 6. 0.3 - 0.8 MeV electrons observed by Helios-1 and -2 during the August 14 - 22, 1979 period [Kunow *et al.*, 1982].
- Figure 7. Same as Figure 6 except for 12.6 - 22.5 MeV protons [Kunow *et al.*, 1982].
- Figure 8. ISEE-3 observations of low energy protons during the August 14 - 22, 1979 period [Domingo and Sanahuja, 1982].
- Figure 9. P78-1 observations of a massive coronal transient in association with the August 14 1979 flare [Poland *et al.*, 1981].
- Figure 10. Velocity data from the San Diego multi-station IPS (Interplanetary Scintillation) equipment. Shown is an increase in velocity present at 3C 273 and not at 3C 298, indicating that the region of enhanced scintillation and velocity associated with the transient on the solar east limb had reached 3C 273 (elongation 43°) but not 3C 298 (elongation 68°) [Jackson, 1982a].
- Figure 11. Position of the region of enhanced scintillation at ~ 1700 UT, 17 August 1979 from the 3.6 hectare IPS array in England. Depicted is the outermost extent of the IPS region of enhanced scintillation in polar coordinates centered on the Sun's position on 17 August 1979. Solar north is at the top, east is to the left. The region of "no information" is caused partially by interference. The general outline of the region of

Figure 21. ISEE-3 measurements of plasma waves associated with the August 18, 1979 flare [Scarf, 1981].

Figure 22. Height of the shock as a function of time using remote sensing of the shock and the drift rate of the type II burst [Cane *et al.*, 1982].

Figure 23. Velocity of the shock as a function of heliocentric distance. The distance is the mean between heliocentric distances of two detections of the shock. [Cane *et al.*, 1982].

Table 1.**Table 1. Characteristics of 14 August 1979 Flare**

He			
Start (UT)			1243
Max (UT)			1244
End (UT)			~1451
Importance			1B
Location			S22, E73
McMath Region			16224
Microwave Radio			
Frequency (GHz)	2.8	9.4	15
Start (UT)	1242	1242.5	1242
Max (UT)	1251.8	1248.7	1248.3
End (UT)	1400	1326.8	-
Type	GB	C	GB
Peak Flux (sfu)	4030	12139	2000
Metric and Decimetric Radio			
Wavelength range	Metric	Decimetric	
Start (UT)	1242.7	1243	
End (UT)	1416	1318	
Intensity	3	3	
Type	IV	IV	
Hard X-rays (ISEE-3)			
Type	Impulsive	Gradual	
Time of max (UT)	1243.1	1248.1	
Peak Flux at ~30 keV	2	12	
(photons cm ⁻² sec ⁻¹ keV ⁻¹)			
Coronal Transient (P78-1)			
Time (UT)			1337
Mass (grams)			~10 ¹⁶
Speed (km sec ⁻¹)			900

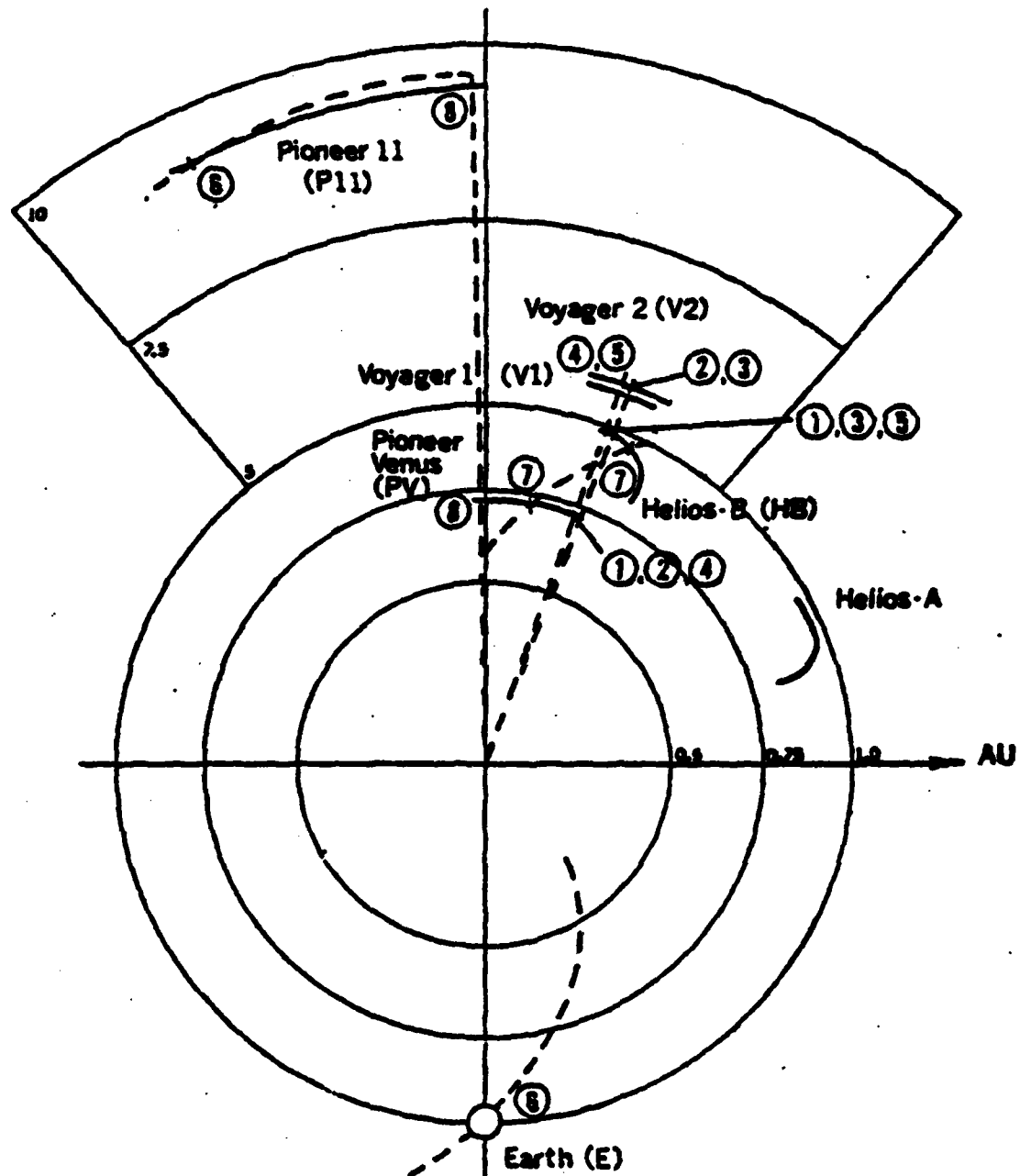
Table 2.

Table 2. Characteristics of 18 August 1979 Flare

Hα		
Start-end (UT)	?	
Location	Behind east limb	
Importance	Large (?)	
Microwave Radio (2.8 GHz)		
Start (UT)	1345.5	
Max (UT)	1423.5	
End (UT)	1450	
Type	C	
Peak Flux (S.F.U.)	490	
Metric and Decimetric Radio		
Wavelength range	Metric	Decimetric
Start (UT)	1409.0	1401.0, 1410.2
End (UT)	1456.0	1403.0, 1411.9
Intensity	3	2, 2
Type	IV	IIIG, IIIG
Kilometric Radio (ISEE-3)		
Start (UT)	~1425	1700
End (UT)	-	1400
Type	III	II
Hard X-rays (ISEE-3)		
Time of max (UT)	1347.5	1412.3
Peak Flux at ~30 keV (photons cm ⁻² sec ⁻¹ keV ⁻¹)	35	140
Energetic Particles (Helios-2)		
Type	Electrons (0.3 - 0.8 MeV)	Protons (27 - 37 MeV)
Start (UT)	~1430	~1430
Max (UT)	~1800	~1800
Peak Flux	-	-
Interplanetary Shock		
Distance from Sun	13.1 R _☉	~1 A.U.
Time (UT)	1500	0552 (20 August)
Speed (km sec ⁻¹)	3500	727
Geomagnetic Effects		
Sudden Commencement	~0600 UT (20 August)	

STIP INTERVAL VII (Start of Solar Maximum Year)

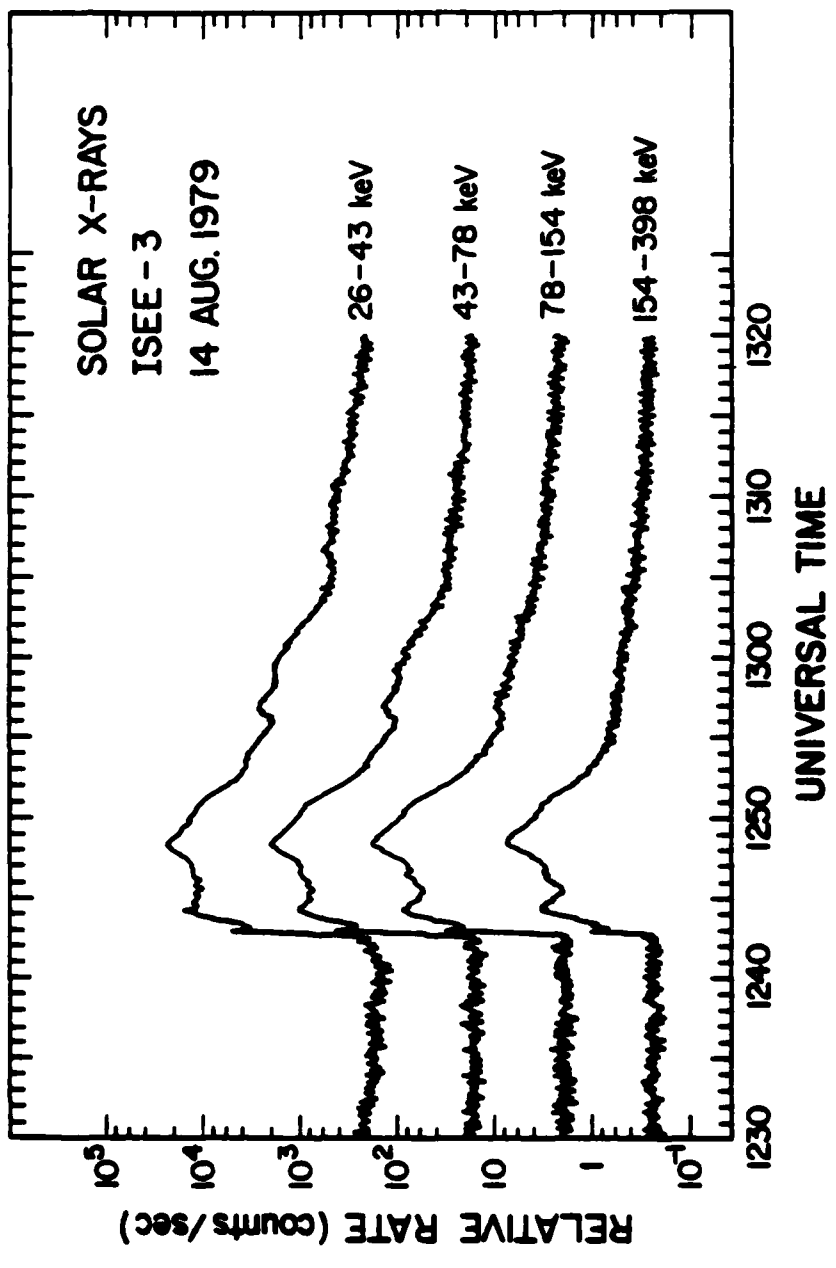
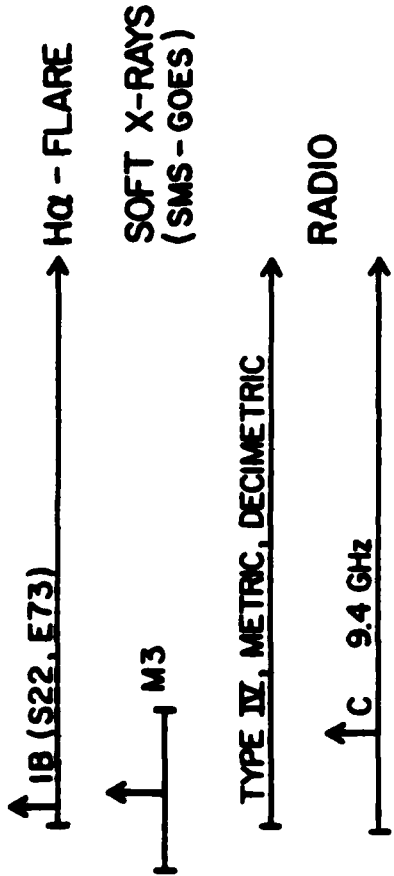
1 AUGUST - 30 SEPTEMBER 1979



Special Conjunctions

1	PV & HB;	RA	8/06	2.0h	5	HB & V2;	RA	8/08	3.5h	
2	PV & V1;	RA	8/06	3.0h	6	E & P11;	IMF	8/13	18.0h	$\tau = 0.0d$
3	HB & V1;	RA	8/06	4.0h	7	PV & HB;	IMF	8/22	15.0h	$\tau = 11.5d$
4	PV & V2;	RA	8/07	0.0h	8	PV & P11;	RA	9/09	14.0h	

Fig. 1



HARD X-RAY SPECTRUM (ISEE-3)

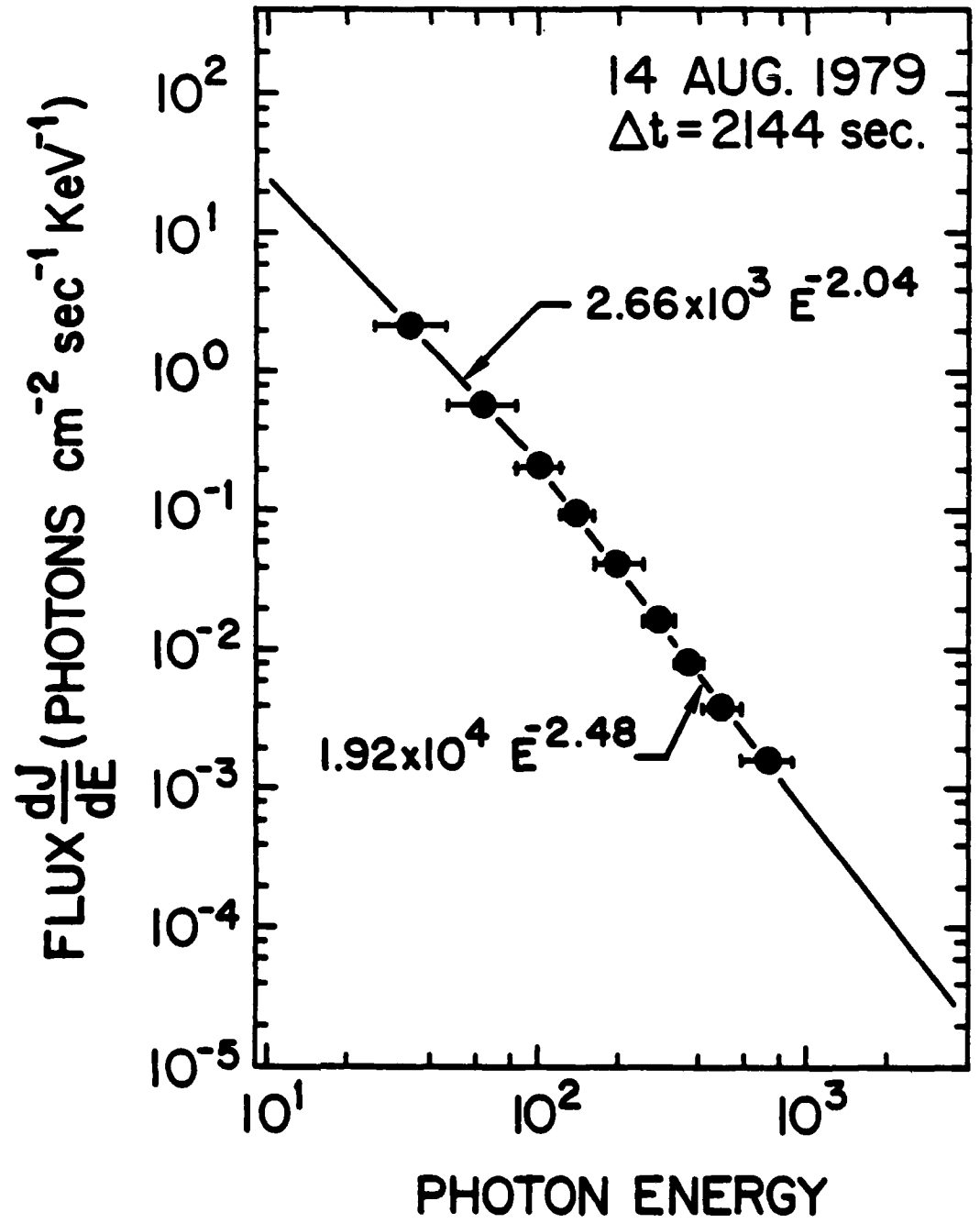


Fig. 3

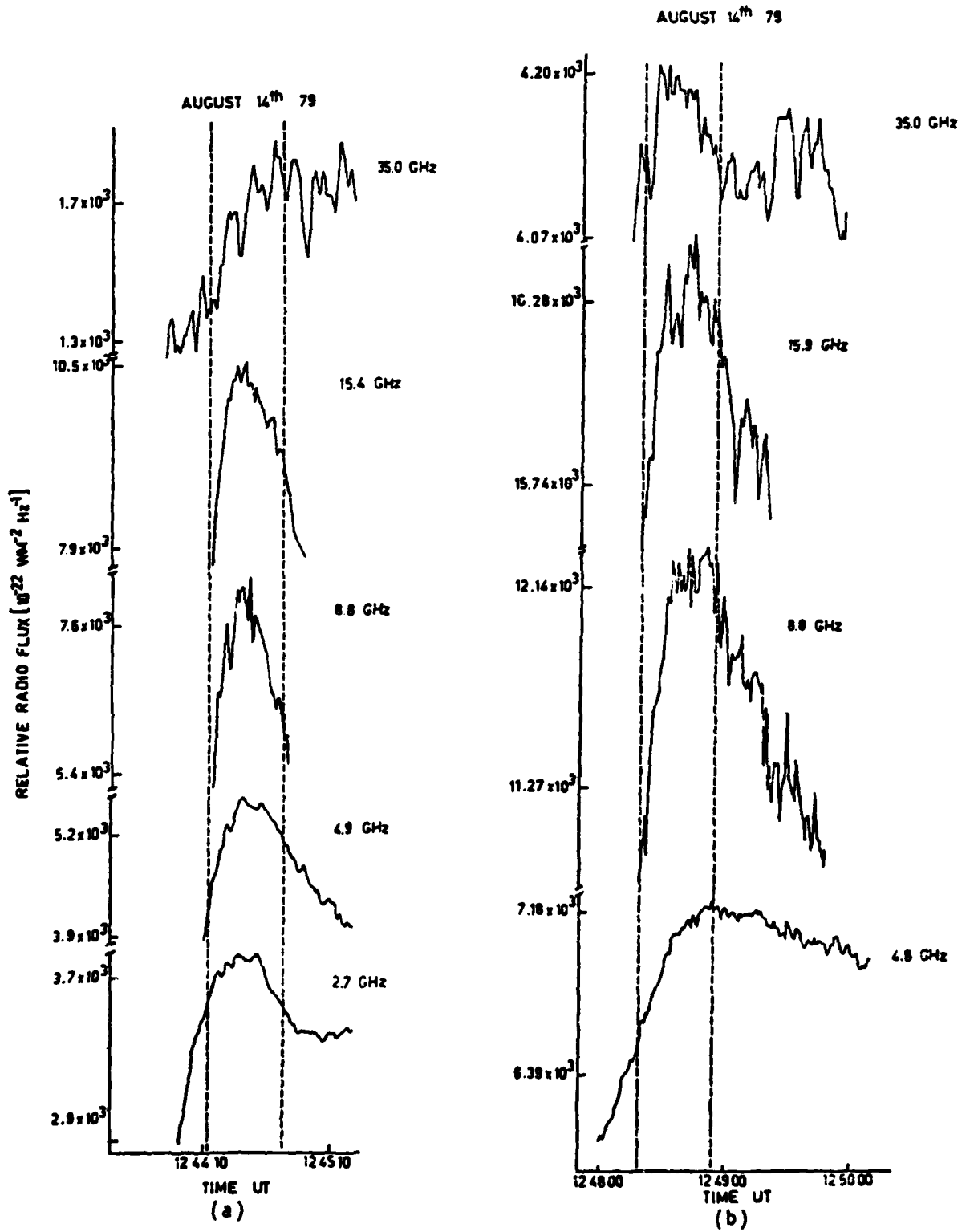


Fig. 4

2.8 GHz RADIO EMISSION
(OTTAWA)

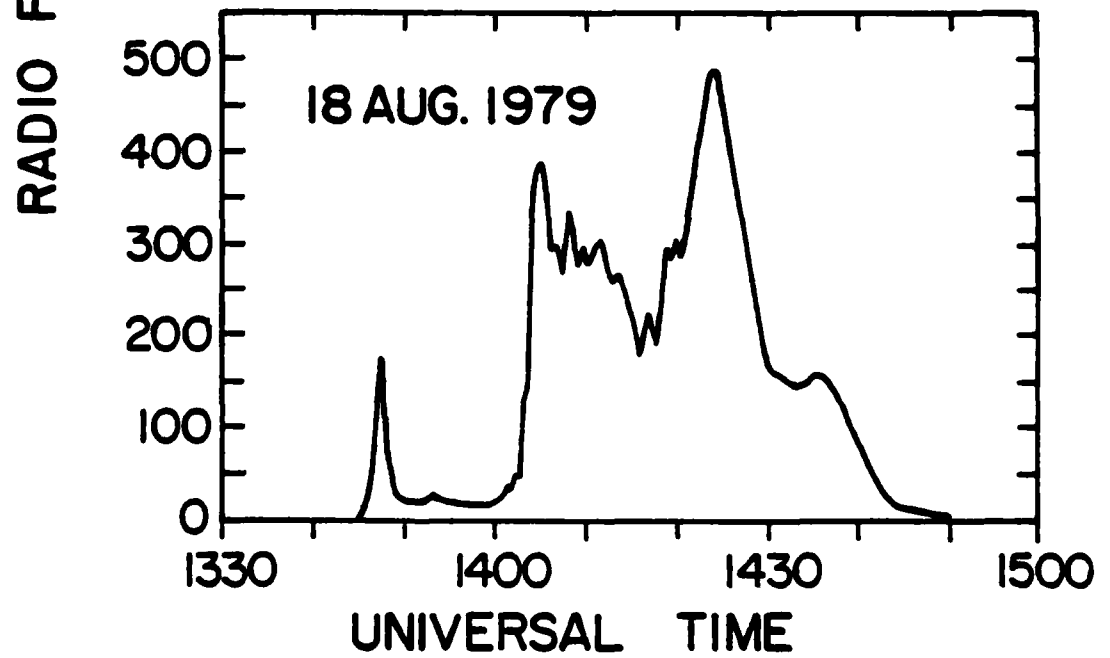
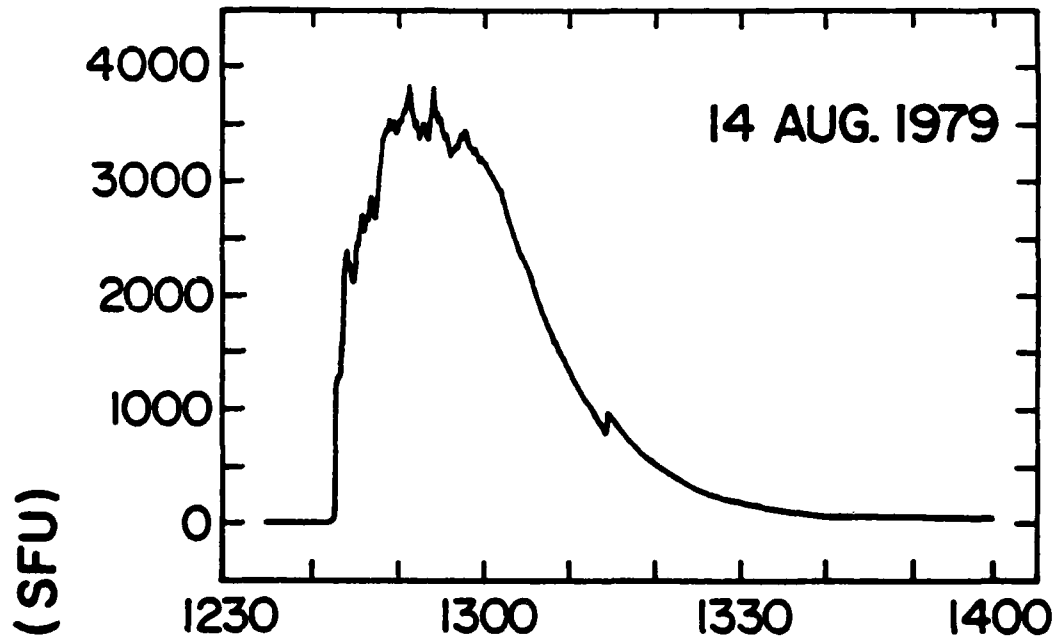


Fig. 5

0.3-0.8 MeV ELECTRONS
(WIBBERENZ et al, 1981)

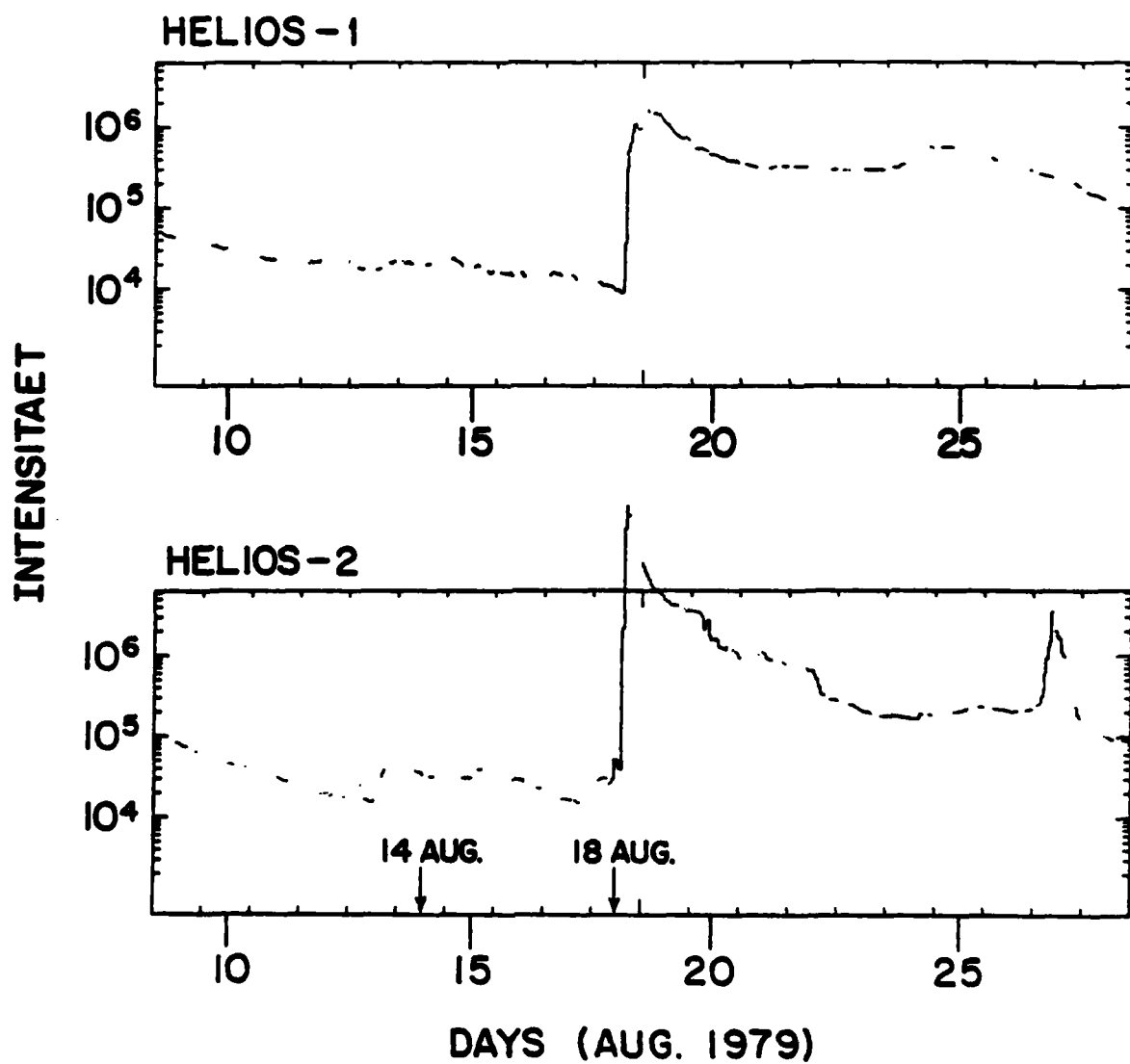


Fig. 6

13 - 27 MeV PROTONS
(WIBBERENZ et al, 1981)

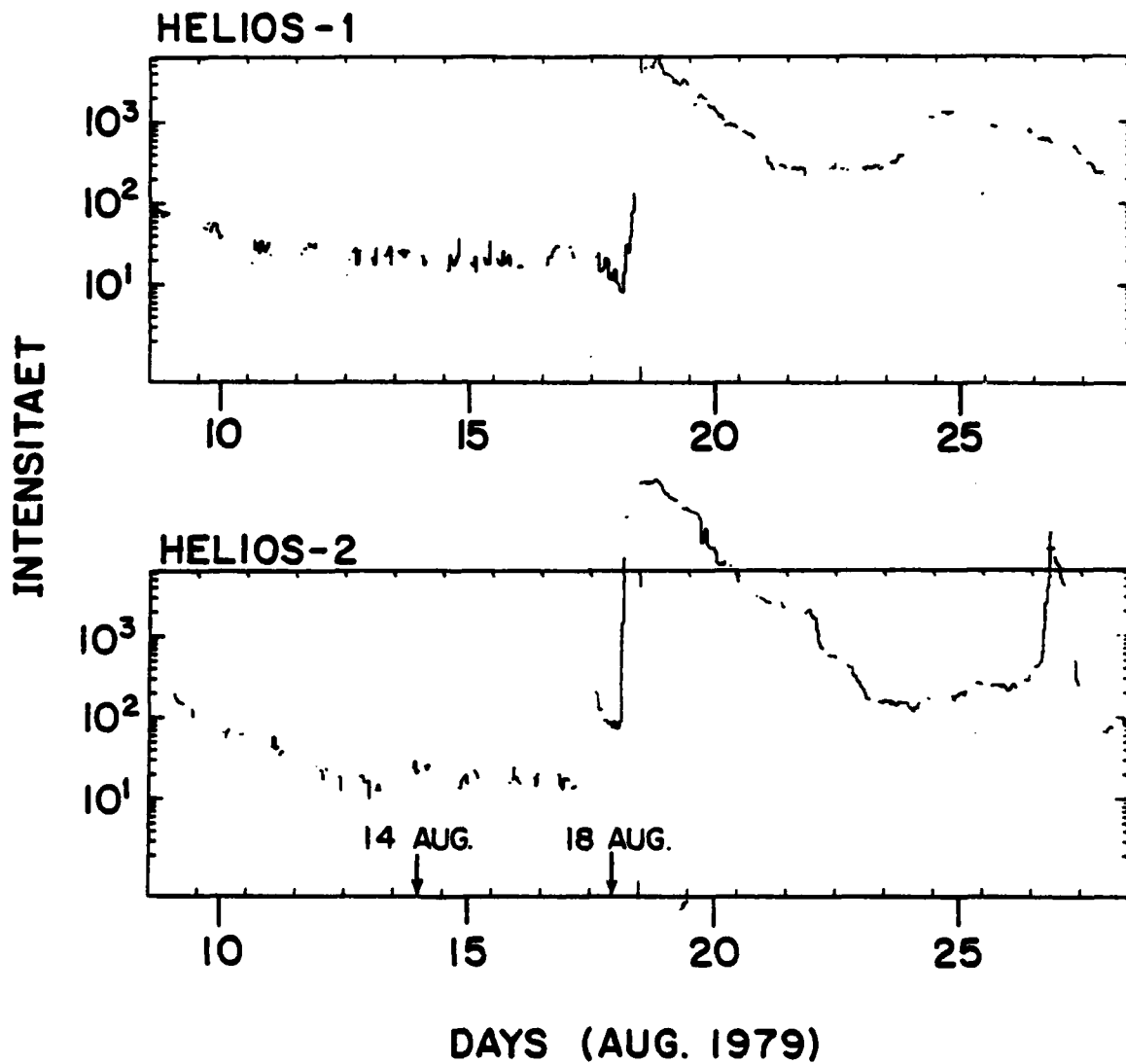
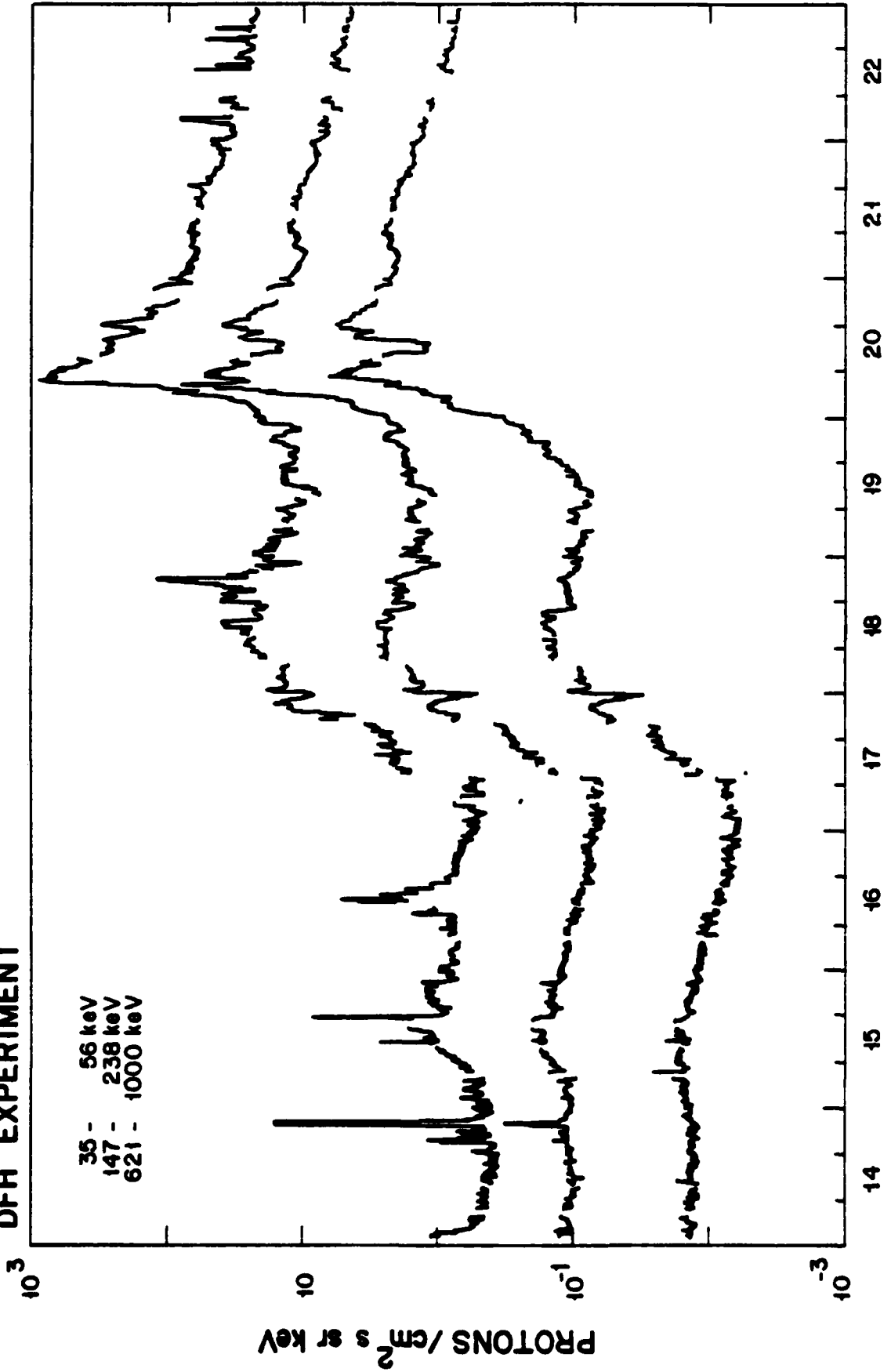


Fig. 7

3/11364/2

ISEE-3 DFH EXPERIMENT

35 - 56 keV
147 - 238 keV
621 - 1000 keV



AUGUST 1979

PRE-EVENT

N



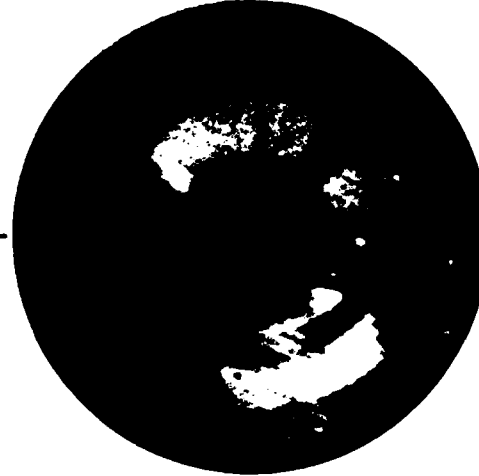
E

1026 UT

14 AUG. '79

1307 UT

N



N

W

1337 UT

|



E

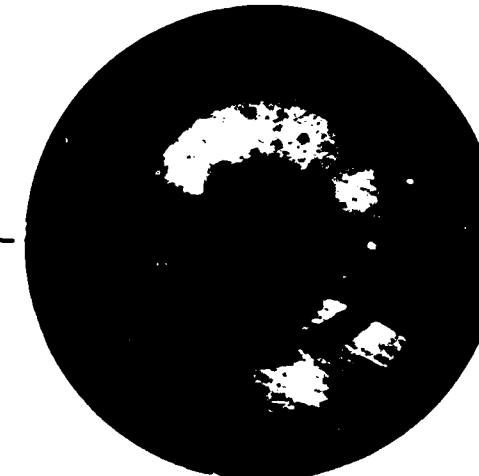
1347 UT

S



1436 UT

S



1456 UT

S

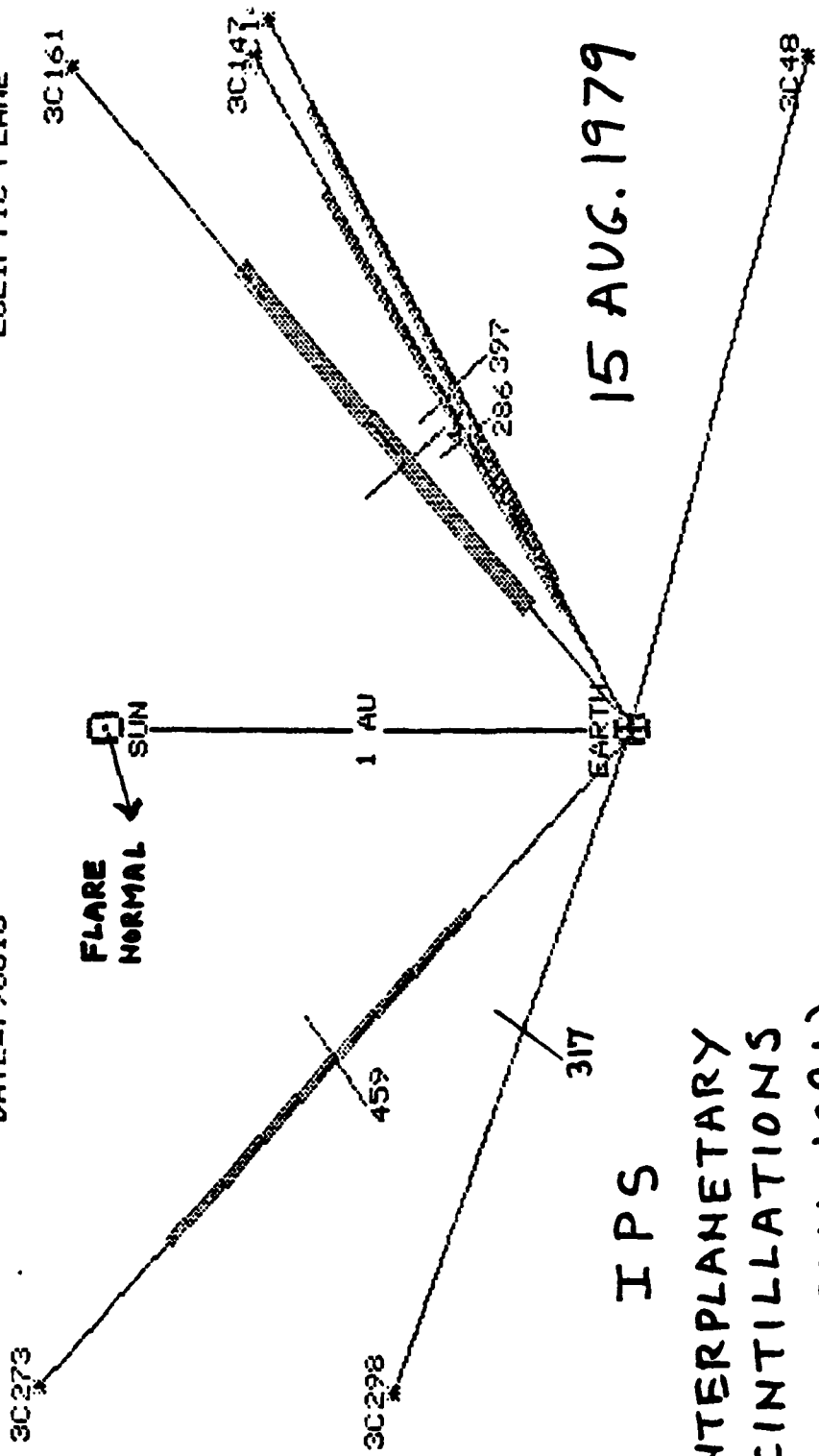
W

HR	KM/S	1-4	DEG	DEG	AU	DEG	DAY	NUMBER
790815	227	12	-	1	3C48	106	13	
	227	14	-	1	3C144-	59	6	
	227	15	286	2	3C147	60	22	0.87 28 1.22 1684.80
	227	16	397	2	3C144	58	6	0.85 32 1.68 1684.85
	227	17	447	3	3C161	52	-15	0.79 32 1.52 1684.88
	227	23	459	2	3C273	43	8	0.68 -47 -4.89 1685.12
790816	228	1	317	2	3C298	68	13	

DATE AND DOY ARE IN UT. VMI IS A MEAN VELOCITY IN THE IPS SCATTERING REGION. THIS REGION'S CENTROID IS SPECIFIED BY HELIOGRAPHIC COORDINATES LAT, DIST, DLON LAG (- TO EAST OF SUN) AND ROTATION NUMBER GIVE MAPPING TO EARTH AND SUN AT VMI

DATE=790816

ECLIPTIC PLANE



IPS
INTERPLANETARY
SCINTILLATIONS
(JACKSON, 1981)

REGION OF IPS ENHANCEMENT

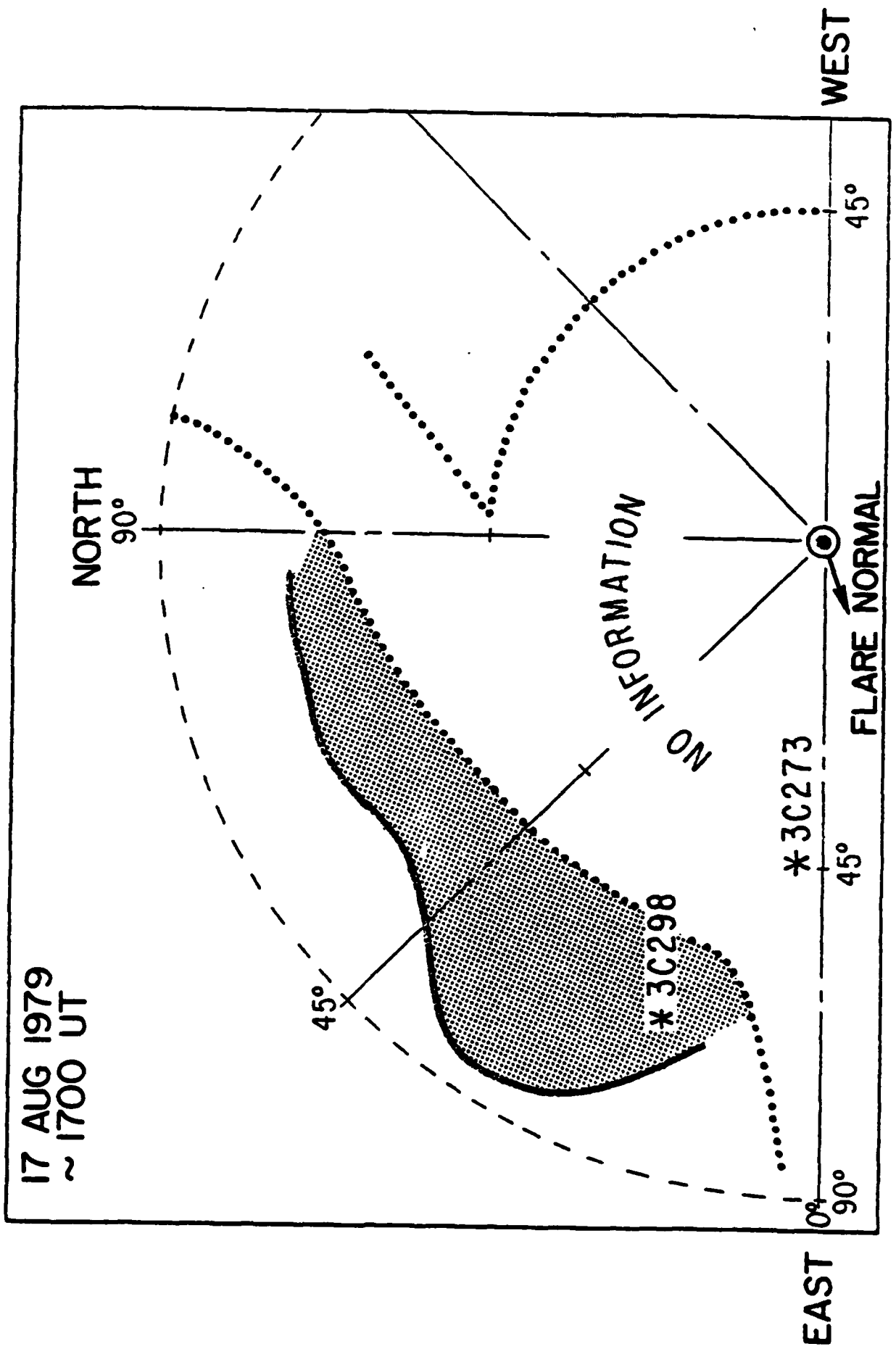


FIG. 11

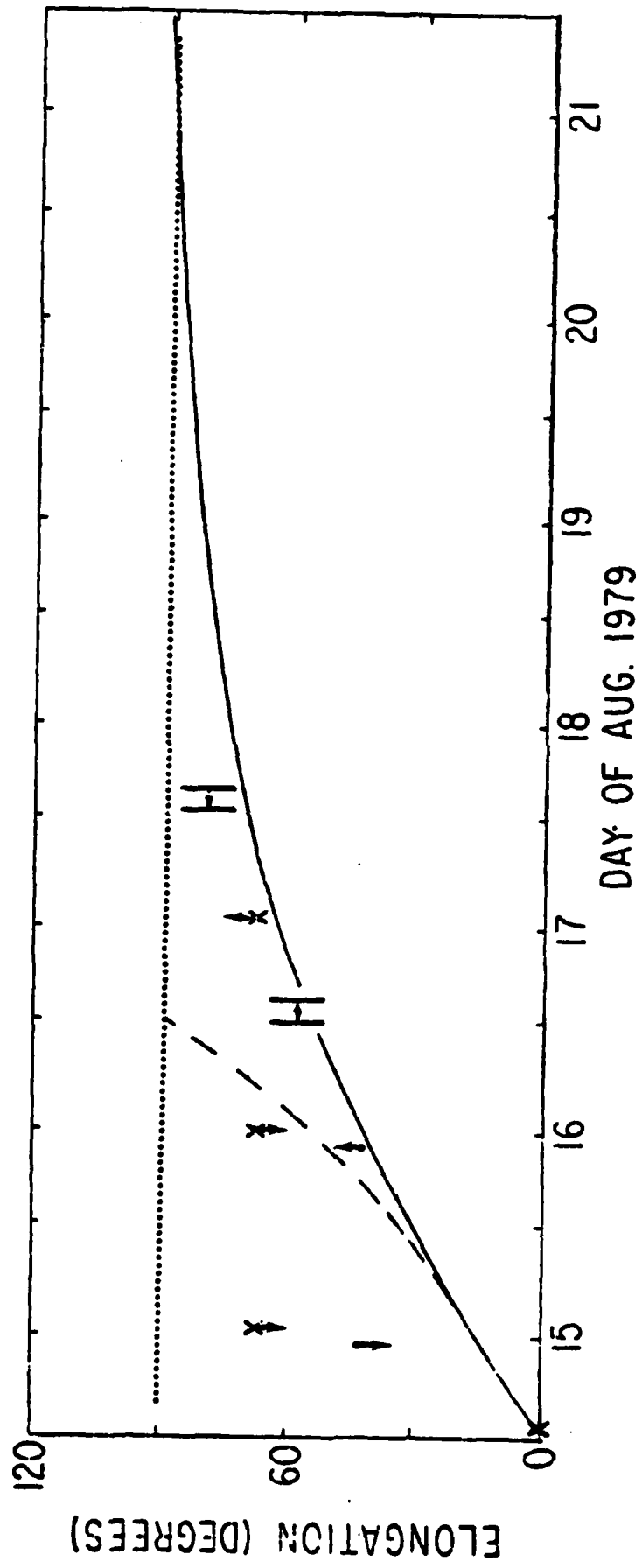


Fig. 12

IMP 8 SOLAR WIND PLASMA

AUGUST 1979

MIT/CSR

IMP 8 PLASMA PARAMETERS

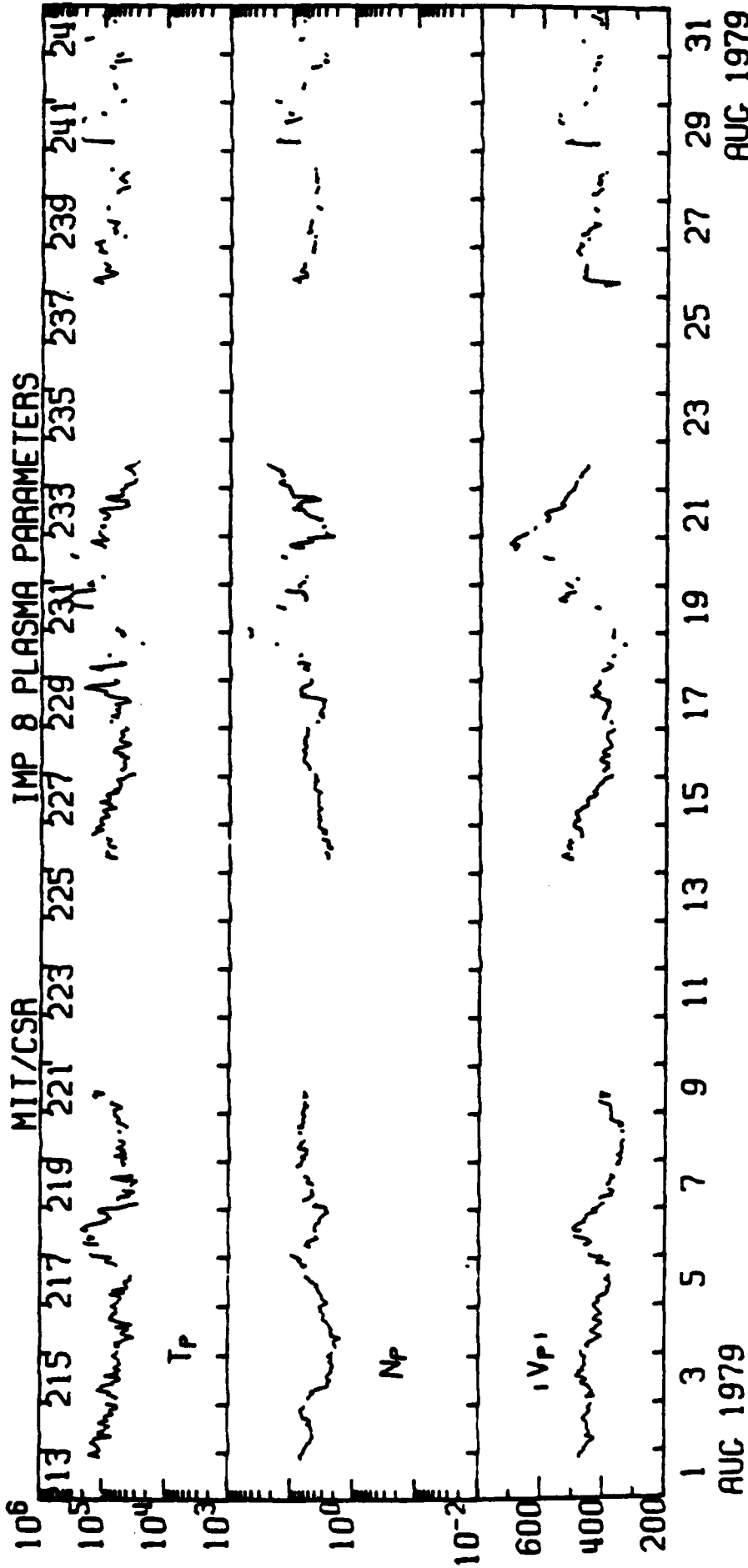


FIG. 13

PIONEER XII

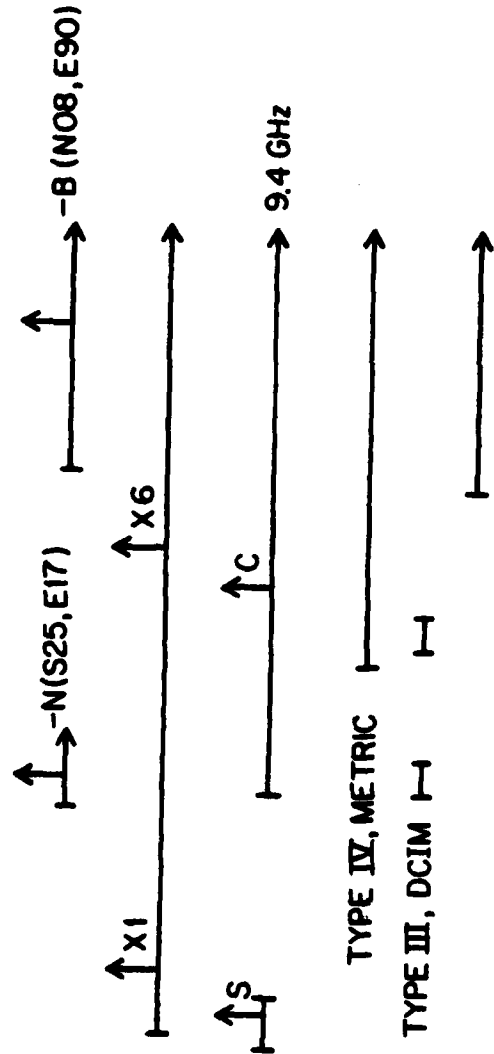
AUGUST 1979

DATE Aug '79	TIME (UT)	ESV (°)	UH+ (Km/sec)	NH+ (H ⁺ /CC)	TH+ (x10 ⁶ °K)
1	1147	164.	428.	22.3	0.203
2	1117		573.	5.3	.436
3	1212		473.	6.8	.13
4	1220		837.	26.3	.519
5	1310		618.	4.8	.282
6	0855		495.	5.4	.078
7	0521		419.	6.6	----
8	1524		323.	13.2	.032
9	1006		318.	26.1	.048
10	1022		313.	56.4	.047
11	1129		471.	41.9	.721
12	0653		562.	8.8	.117
13	0000		528.	12.9	.279
14	2207		597.	4.4	.102
15	2358	173.	459.	----	----
16	0007		----	10.8	----
17					
18					
19					
20					
21					
22					
23					
24					
25					
26					
27					
28					
29					
30		176.			
31					

NOTE: PN-12 data estimated suspended until approximately 9 Sep 79.

Venus undergoing superior conjunction. Communication link immersed in solar corona resulting in somewhat noisy and unsatisfactory data.

Additionally, PN-11 undergoing Saturn encounter thus reducing PN-12 Deep Space Net (DSN) priority.

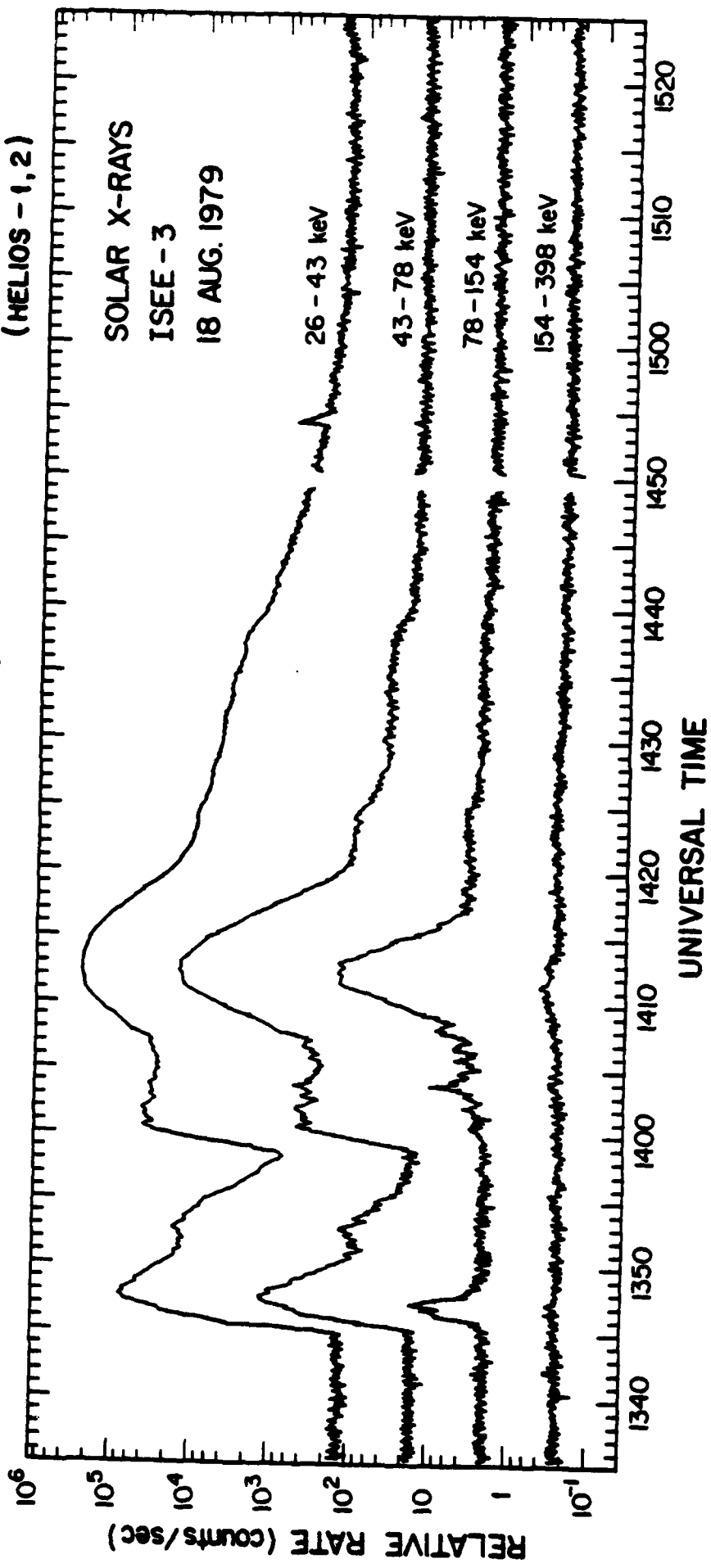


H α - FLARE
 SOFT X-RAYS
 (SMS - GOES)

RADIO

ENERGETIC PARTICLES
 (HELIOS - 1, 2)

SOLAR X-RAYS
 ISEE - 3
 18 AUG. 1979



HARD X-RAY SPECTRUM (ISEE-3)

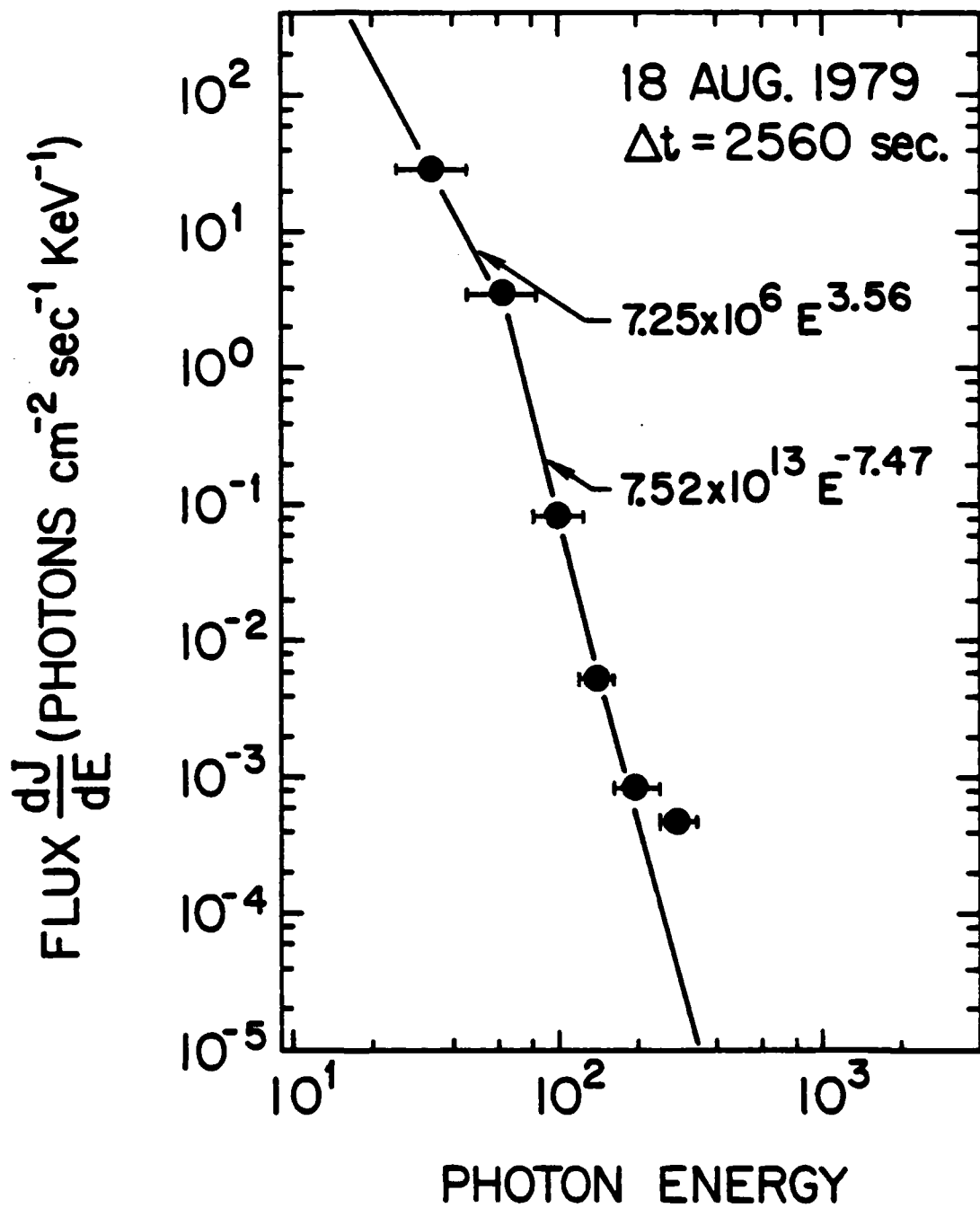


Fig. 16

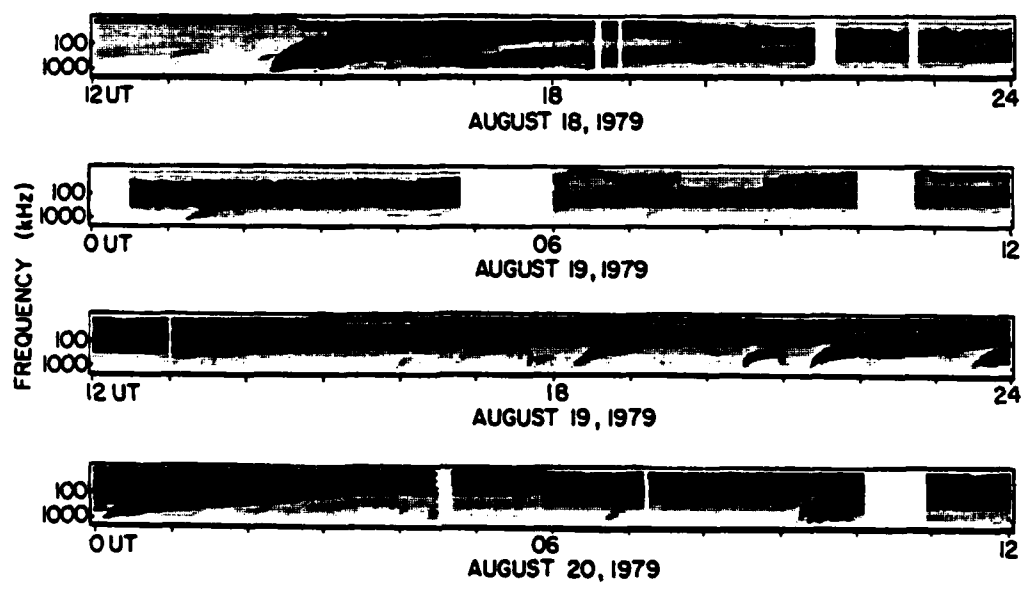
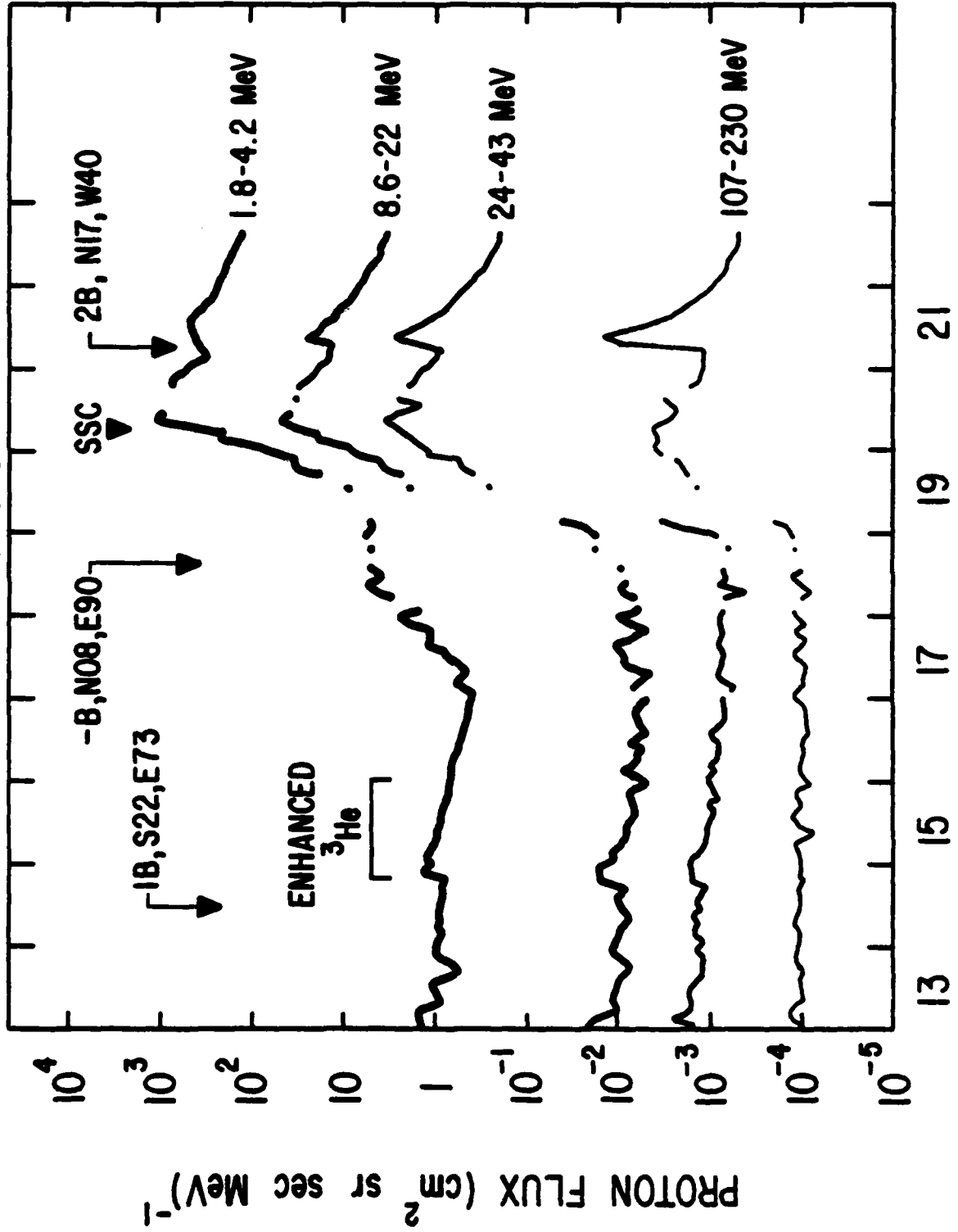


Fig. 17

IMP-8 / GSFC



AUGUST 1979

SPECTRUM AT TIME
OF MAXIMUM FLUX

19-20 AUG 1979 IMP-8/GSFC

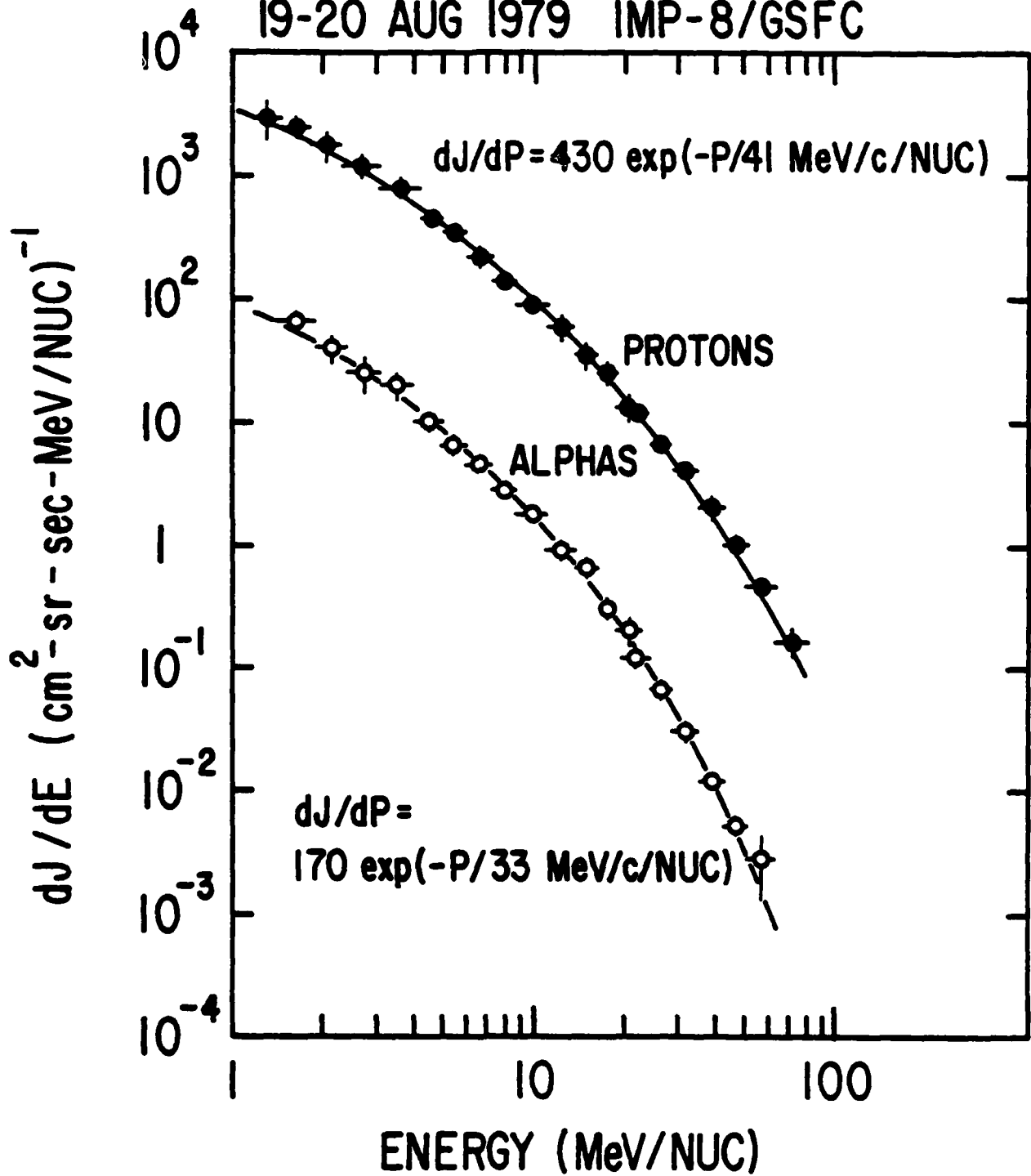
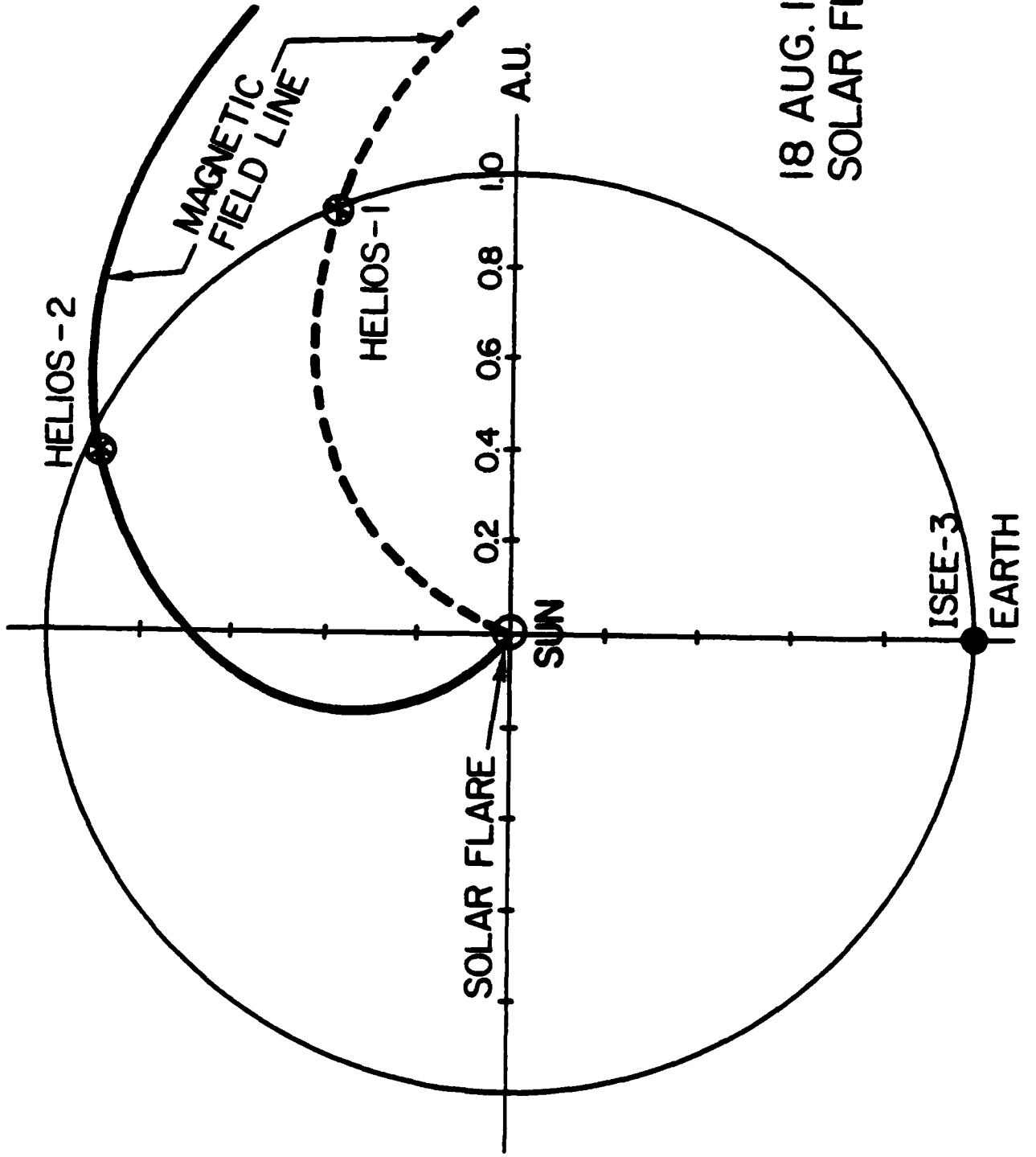


Fig. 18b

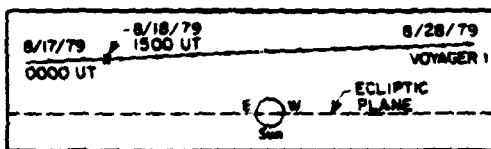


18 AUG. 1979
SOLAR FLARE

Fig. 19

SHOCK WAVE INFERRED FROM RADIO MEASUREMENTS
18 AUG. 1979 FLARE
(WOO & ARMSTRONG, 1981)

VIEW FROM
THE EARTH



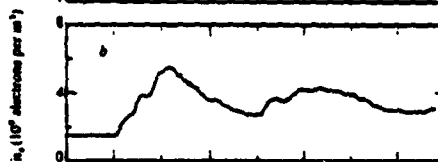
SHOCK AT 13.1 R_{\odot}

← COMPRESSED REGION | DRIVER GAS →

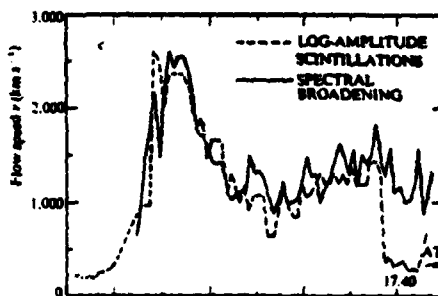
TURBULENCE
PARAMETERS



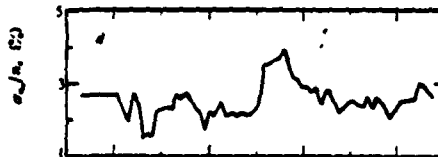
ELECTRON
DENSITY



SOLAR WIND
FLOW SPEED



R.M.S. ELECTRON
DENSITY
FLUCTUATIONS



UNIVERSAL TIME

Fig. 20

PLASMA WAVES (ISEE-3)
(SCARF, 1981)

SHOCK ↓

20 AUG. 1979

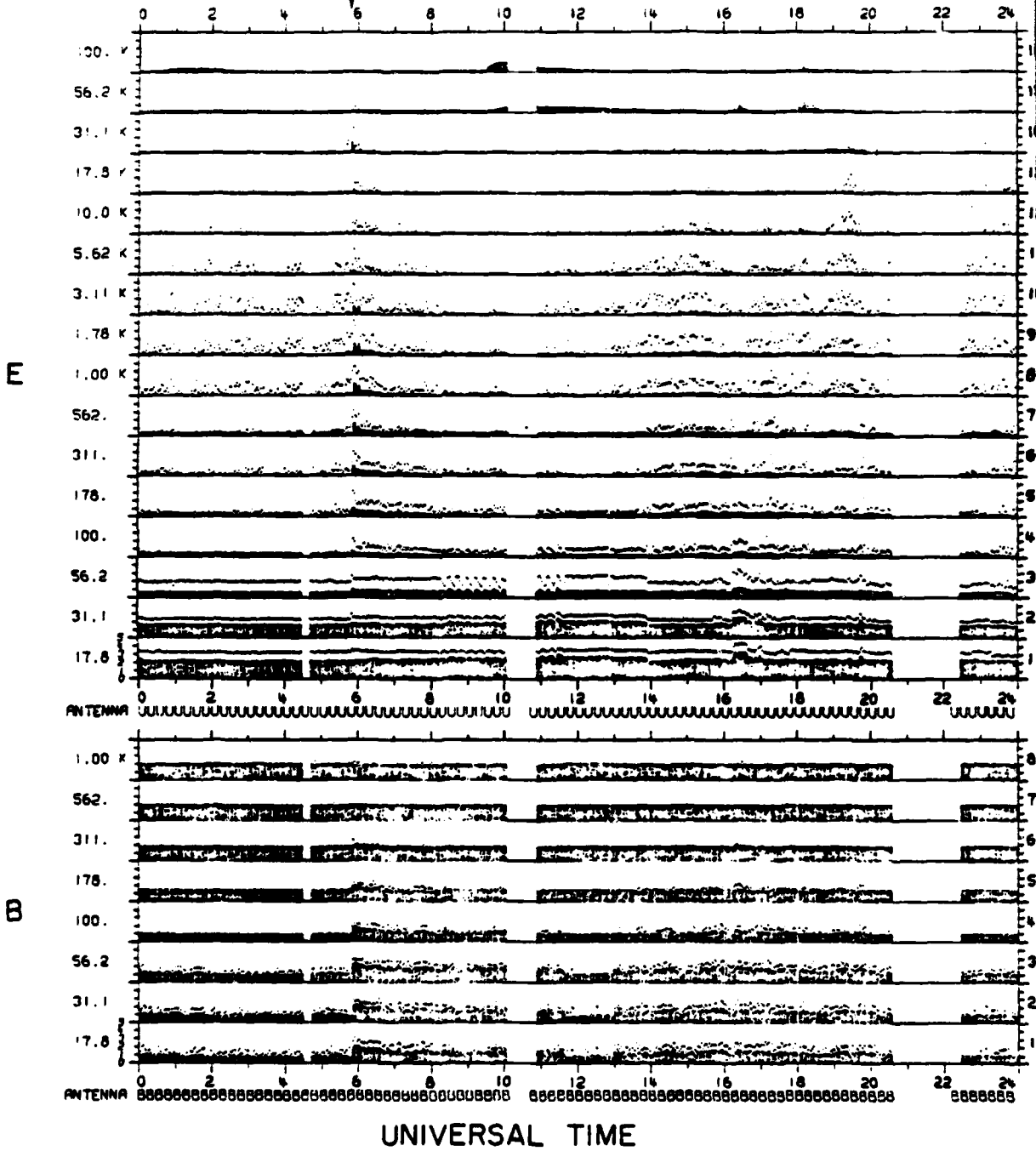
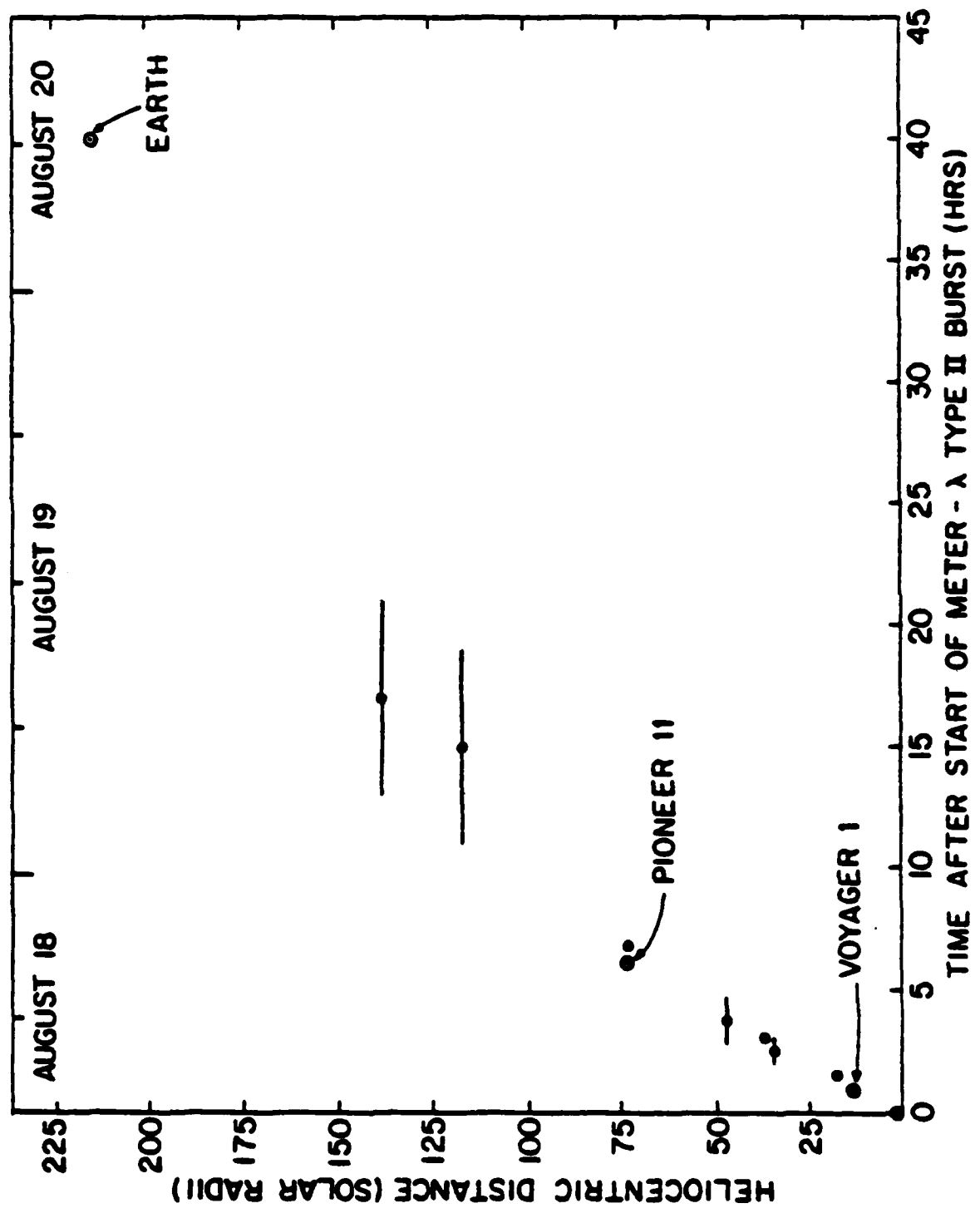


Fig. 21



END

FILMED

3-83

DTIC